Проект SATURNE: общий обзор

К.А. Кузаков

МГУ имени М.В. Ломоносова

The **Sarov Tritium Neutrino Experiment (SATURNE)** is part of the research program of the **National Center for Physics and Mathematics** founded in 2021

The main goals of SATURNE are

- **first observation of coherent elastic neutrino-atom scattering (CEνAS)**
- **search for neutrino magnetic moment with neutrino-atom scattering**

using a high-intensity tritium neutrino source: at least **1 kg, possibly up to 4 kg of T_2**

CEνAS vs CEνNS

CEνAS: Coherent Elastic Neutrino-Atom Scattering

predicted by *Yu. V. Gaponov and V. N. Tikhonov, Yad. Fiz. (USSR)* ***26*** *(1977) 594 (in Russian);* ***no experimental observation so far***

CEνNS

- $|\vec{q}| R_{\text{nuc}} \ll 1$

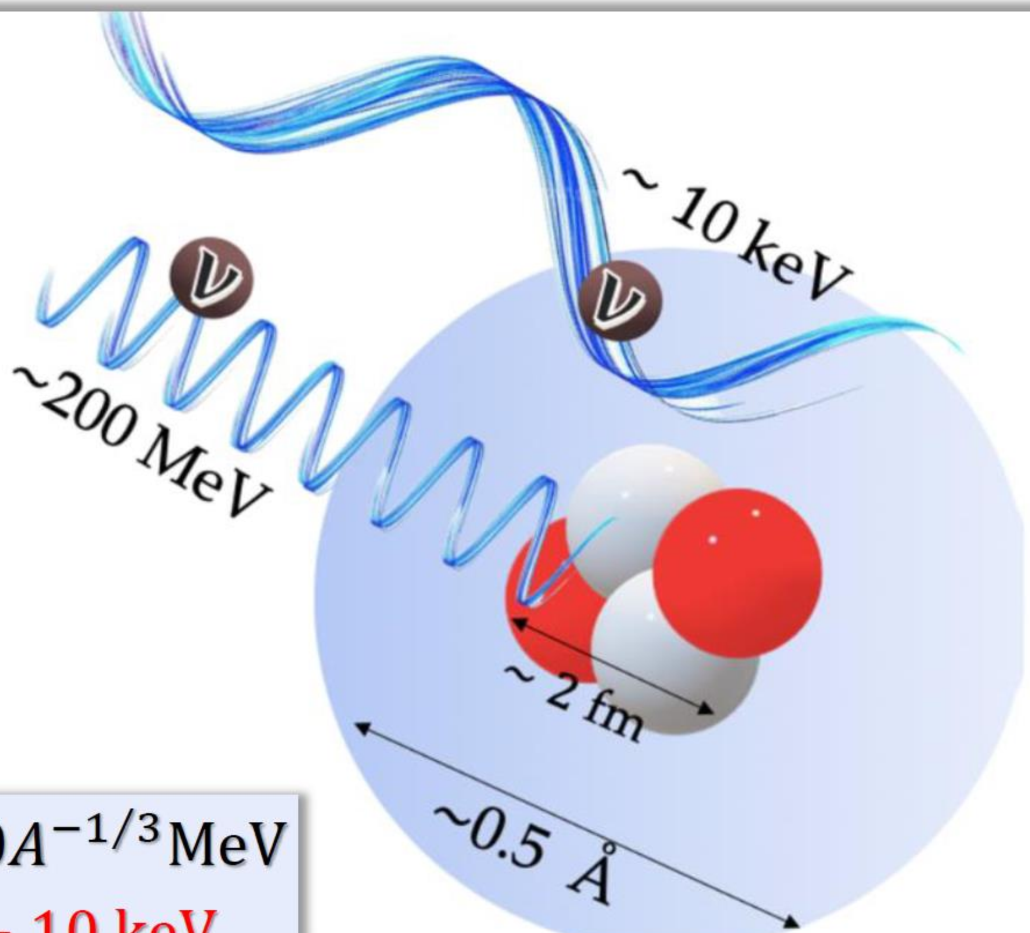
$\vec{q}$ is the momentum transfer

R_{nuc} is the nuclear radius

CEνAS

- $|\vec{q}| R_{\text{atom}} \ll 1$

R_{atom} is the atomic radius

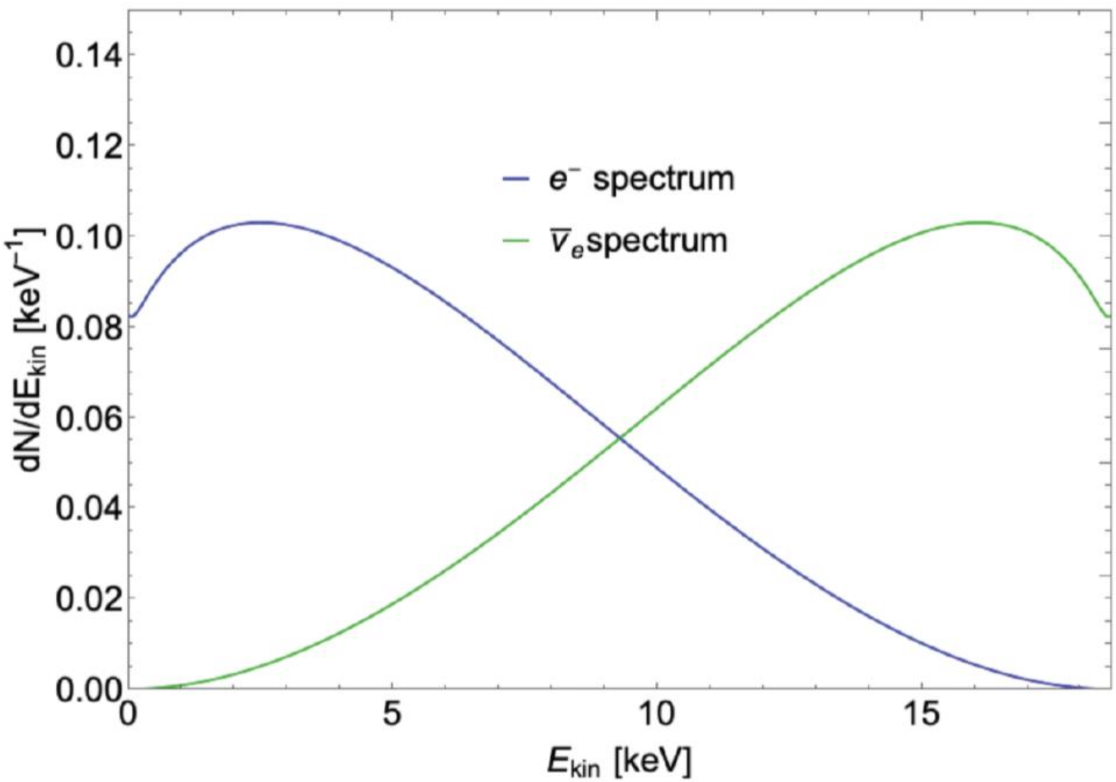


CEνNS: $E_\nu \lesssim 1/R_{\text{nuc}} \sim 200 A^{-1/3}\,\text{MeV}$

CEνAS: $E_\nu \lesssim 1/R_{\text{atom}} \sim 1 - 10\ \text{keV}$

Tritium neutrinos

$$^3\mathrm{H} \rightarrow \ ^3\mathrm{He} + e^- + \bar{\nu}_e$$



$$Q = 18.6\ \mathrm{keV}$$

$$t_{1/2} = 12.3\ \mathrm{yrs}$$

$$\langle E_{\bar{\nu}_e} \rangle = 12.9\ \mathrm{keV}$$

With 1-4 kg of tritium the neutrino flux in SATURNE will be $\Phi_{\bar{\nu}_e}$~10^{13}-10^{14} $sm^{-2}s^{-1}$

In contrast to stopped-pion beams ($\langle E_\nu \rangle \sim 30$ MeV) and nuclear reactors ($\langle E_\nu \rangle \sim 1$ MeV), **with a tritium neutrino source it is possible to fulfill the coherence condition in elastic neutrino-atom scattering**

Atomic recoil energy scale in CEνAS

From conservation of energy and momentum:

$$T_R \leq \frac{2E_\nu^2}{m}$$

T_R is atomic recoil energy

$m \approx A$ GeV is atomic mass

In the reactor CEνNS experiment CONNIE: Threshold is 15 eV$_{ee}$ (with CCD sensors)

Aguilar-Arevalo et al., arXiv:2403.15976v1 [hep-ex]

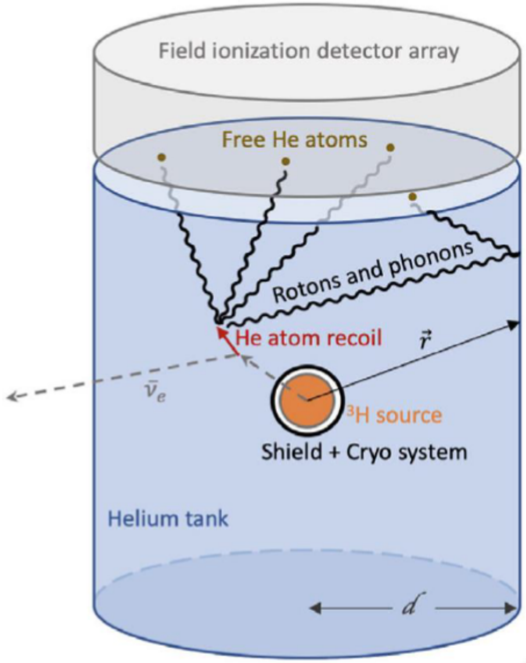
If E_ν~10 keV: $T_R \lesssim \frac{200}{A}$ meV

For the lightest atom (A=1): $T_R \lesssim 200$ meV

Light atomic targets, such as H or He, and new detector technologies are needed to observe CEνAS

Potential of a low-energy detector based on ^{4}He evaporation

M. Cadeddu, F. Dordei, C. Giunti, K. Kouzakov, E. Picciau, A. Studenikin, PRD ***100*** *(2019) 073014*

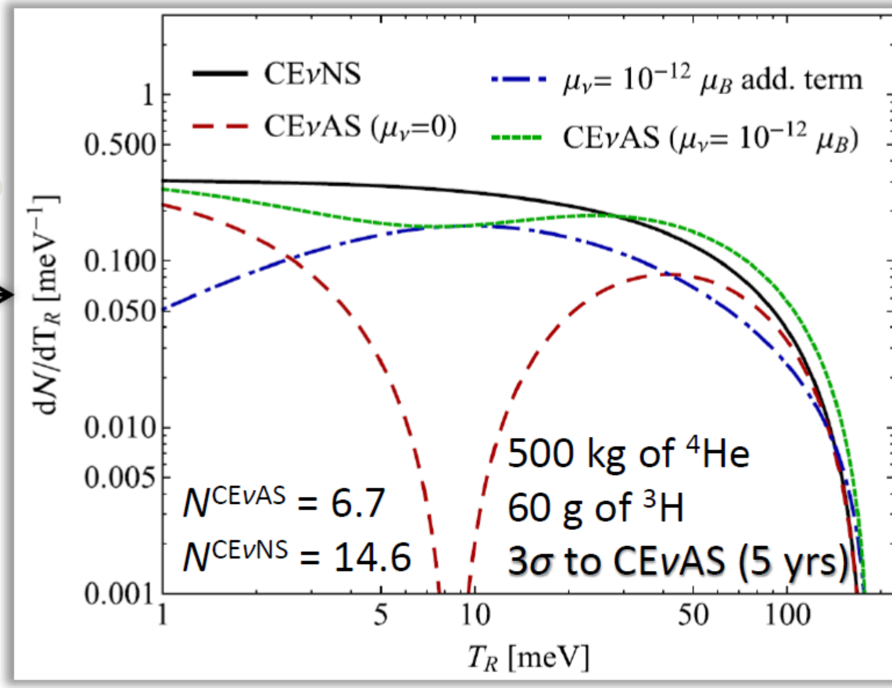


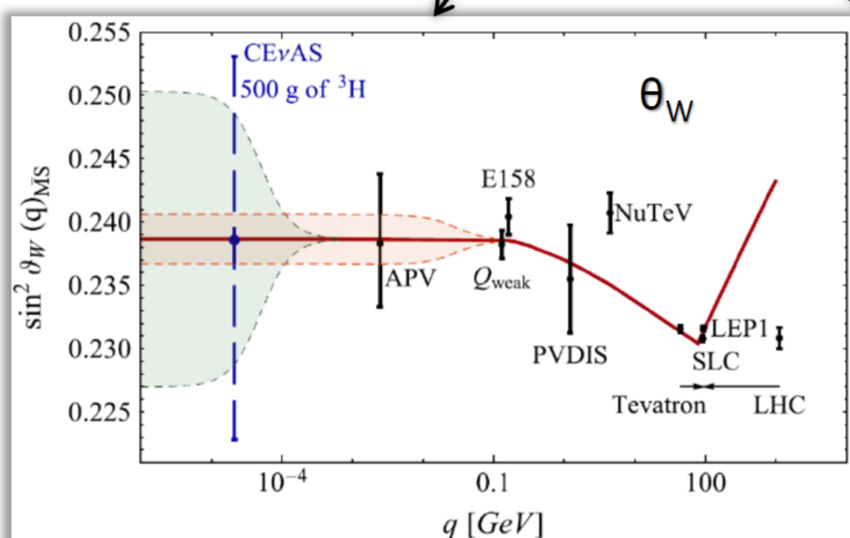
We can study:

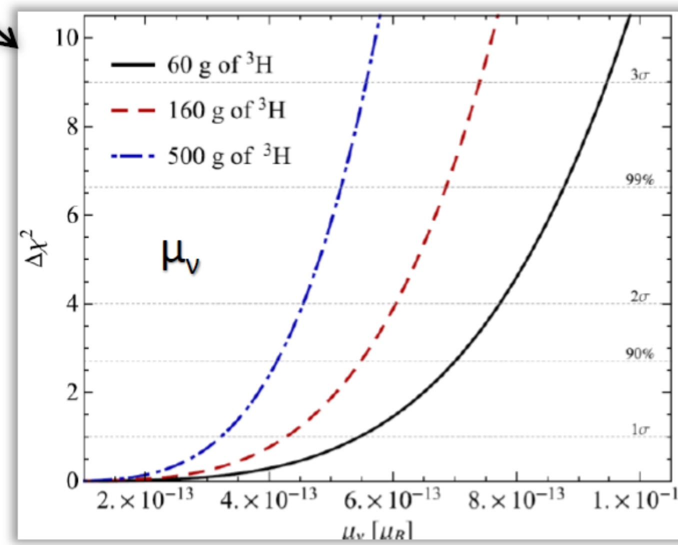
CEνAS

θ_W

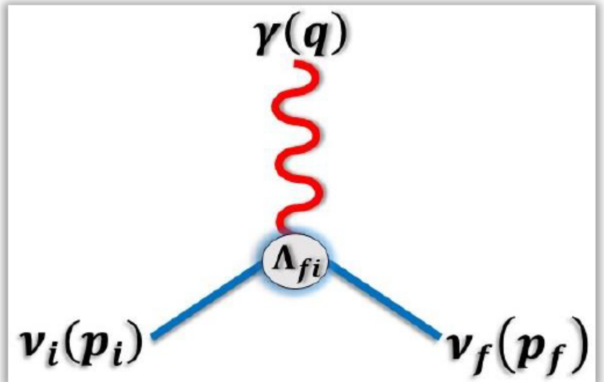
μ_ν







Neutrino magnetic moment



*C. Giunti and A. Studenikin, **Neutrino electromagnetic interactions: A window to new physics**, Rev. Mod. Phys. **87** (2015) 531; arXiv:1403.6344 [hep-ph]*

*C. Giunti, K. Kouzakov, Y.-F. Li, and A. Studenikin, **Neutrino electromagnetic properties**, Annu. Rev. Nucl. Part. Sci. **75** (2025); arXiv:2411.03122 [hep-ph]*

The effective neutrino electromagnetic vertex under the Lorentz and gauge invariance:

$$\Lambda_{\mu}^{(\mathrm{EM};\nu)fi}(q) = \left(\gamma_\mu - \frac{q_\mu q}{q^2}\right)\left[f_Q^{fi}(q^2) - q^2\, f_A^{fi}(q^2)\gamma_5\right] - i\sigma_{\mu\nu}q^\nu\left[f_M^{fi}(q^2) + if_E^{fi}(q^2)\gamma_5\right]$$

In the minimally extended SM with addition of right-handed massive Dirac neutrinos:

$$\mu_\nu \simeq 3.2 \times 10^{-19}\mu_B\left(\frac{m_\nu}{1\ \mathrm{eV}}\right)$$

*K. Fujikawa and R. Shrock, PRL **45** (1980) 963*

$m_\nu < 0.45$ eV at 90% CL

*M. Aker et al. (The KATRIN Collaboration), Science **388** (2025) 180; arxiv:2406.13516*

Much greater μ_ν values are predicted beyond the minimally extended SM

World leading upper bounds on μ_ν

Laboratory bounds (elastic $\nu - e^-$ scattering)

solar neutrinos (XENONnT)

*A. Khan, Phys. Lett. B **837** (2023) 137650*

$\mu_\nu < 6.3 \times 10^{-12} \mu_B$

reactor neutrinos (GEMMA)

*A. Beda et al., Adv. High Energy Phys. **2012** (2012) 350150*

$\mu_{\nu_e} < 2.9 \times 10^{-11} \mu_B$

CEνNS bounds

*V. De Romeri et al., JHEP **04** (2023) 035*

$\mu_{\nu_e} < 3.8 \times 10^{-9} \mu_B$

$\mu_{\nu_\mu} < 2.6 \times 10^{-9} \mu_B$

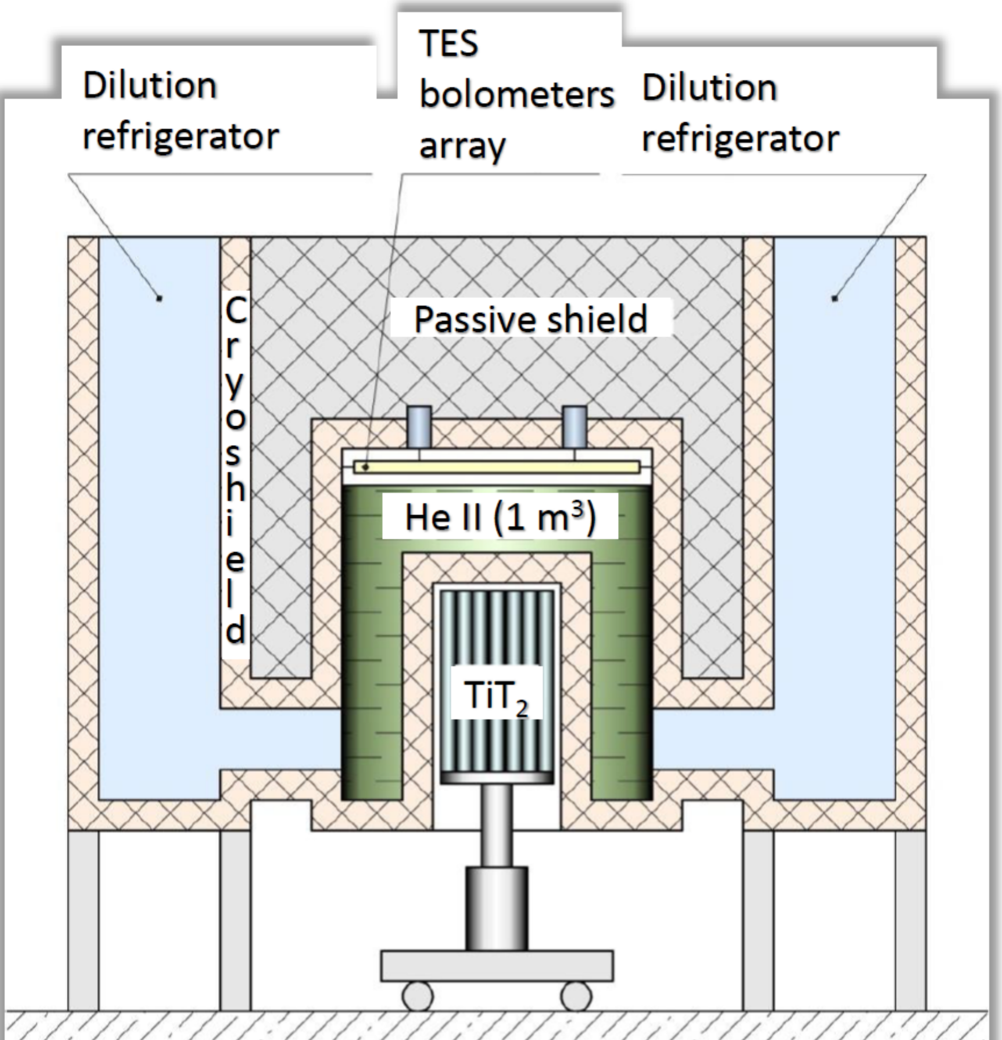
Astrophysical bounds (luminosity of globular star clusters)

*N. Viaux et al., Astron. & Astrophys. **558** (2013) A12; S. Arceo-Diaz et al, Astropart. Phys. **70** (2015) 1; F. Capozzi and G. Raffelt, Phys. Rev. D **102** (2020) 083007*

$\mu_\nu < (1.2-2.6) \times 10^{-12} \mu_B$

With CEνAS, we could improve the CEνNS limits by four orders of magnitude, and the world leading limits by an order of magnitude

He II detector concept to study CEνAS



Helium II detector (1000 L)

- Liquid He-4 at 40-60 mK
- Array of 1000 TESs (transition edge sensors)
- 1000-channel SQUID readout

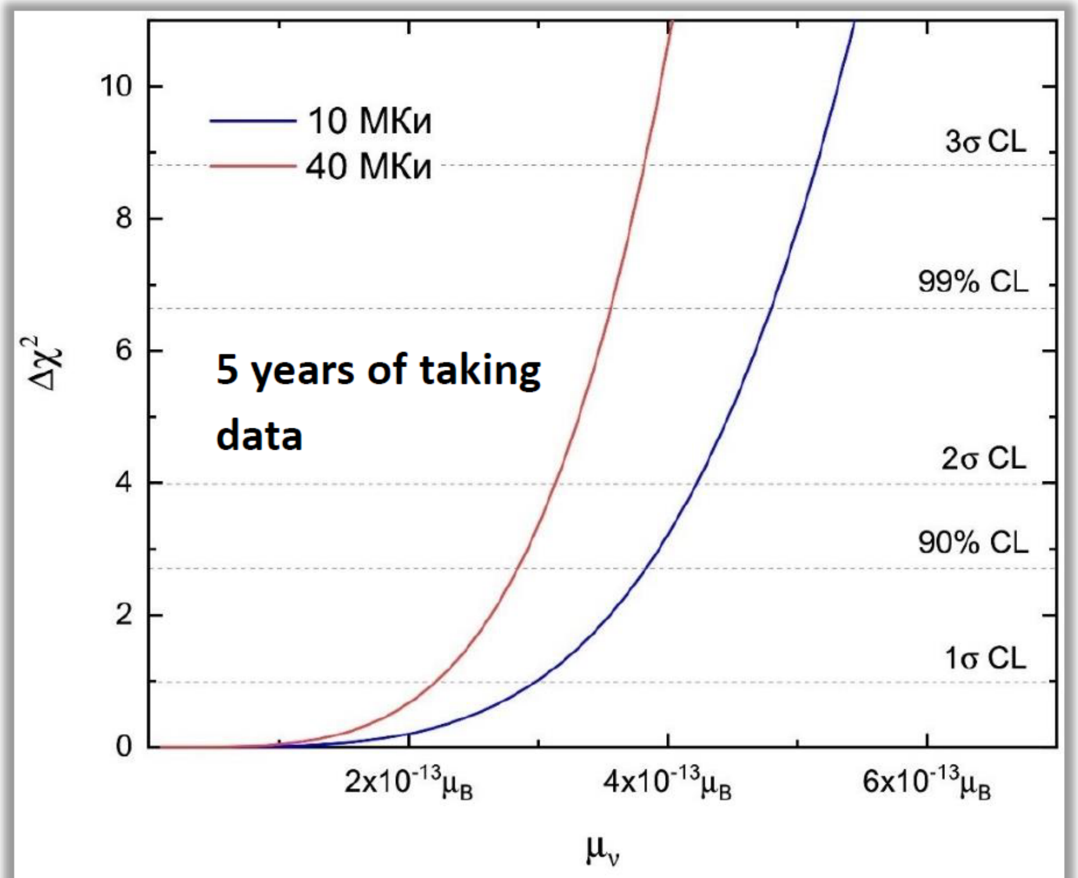
Tritium neutrino source

1-4 kg, 10-40 MCi

- Tubular elements with TiT_2

*A.A. Yukhimchuk et al. Fusion Science & Technology **48**, No.1 (2005) 731-736*

Projected μ_ν-sensitivity of He-4 detector



Tritium mass is

(i) 1 kg (10 MCi)

(ii) 4 kg (40 MCi)

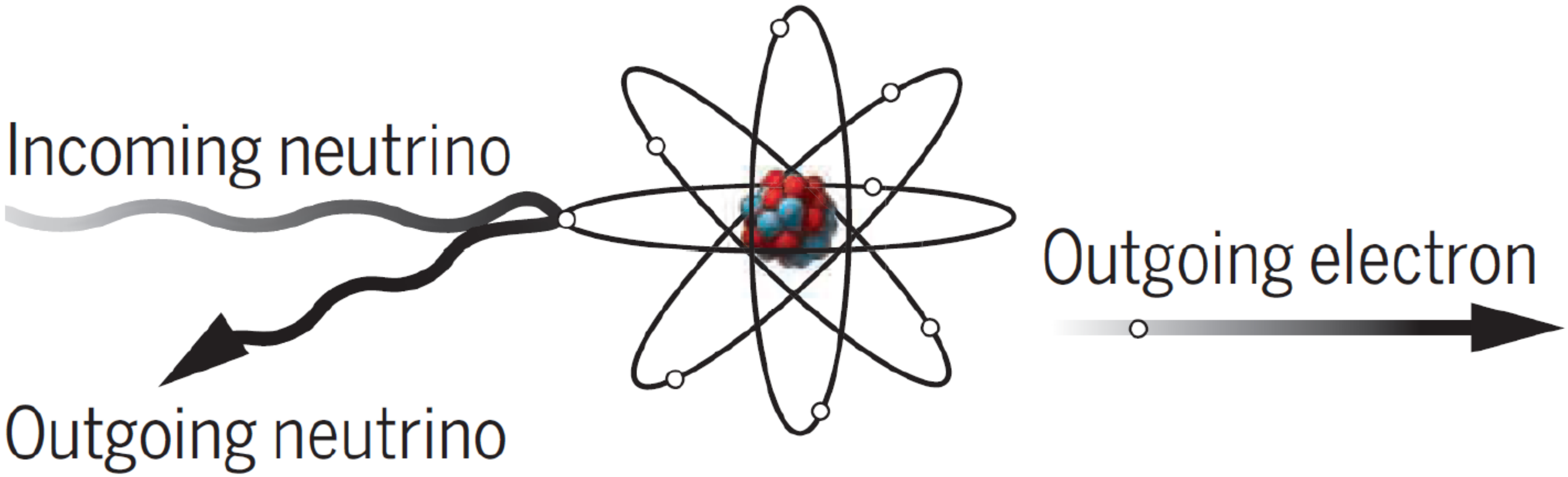
$$\Delta\chi^2 = \chi^2 - \chi^2_{\min}$$

$$\chi^2 = \left(\frac{N_{SM} - N}{\sqrt{N_{SM}}}\right)^2$$

$$N = N_{SM} + N_{\mu_\nu}$$

Initial tritium activity	N_{SM}	N_{μ_ν}, $3 \times 10^{-13}\mu_B$	N_{μ_ν}, $10^{-12}\mu_B$
10 MCi	53.7	7.1	82.1
40 MCi	177.1	24.6	271.1

Atomic ionization channel



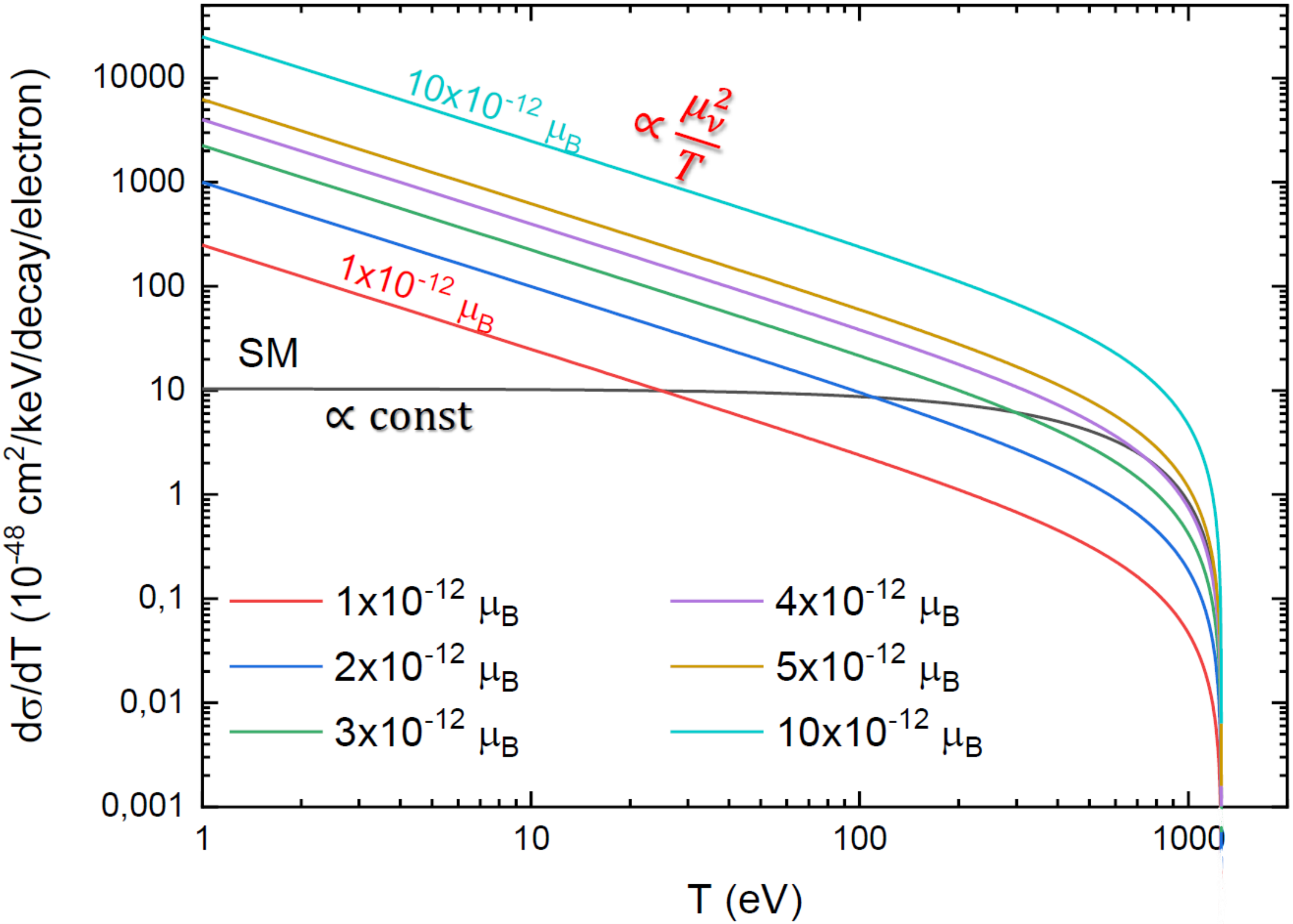
World leading laboratory constraints on μ_ν, like those from XENNONnT and GEMMA, are obtained by studying the atomic ionization channel (elastic $\nu - e^-$ scattering)

In **SATURNE** we develop

- Si crystal detector
- CsI(pure) scintillation detector

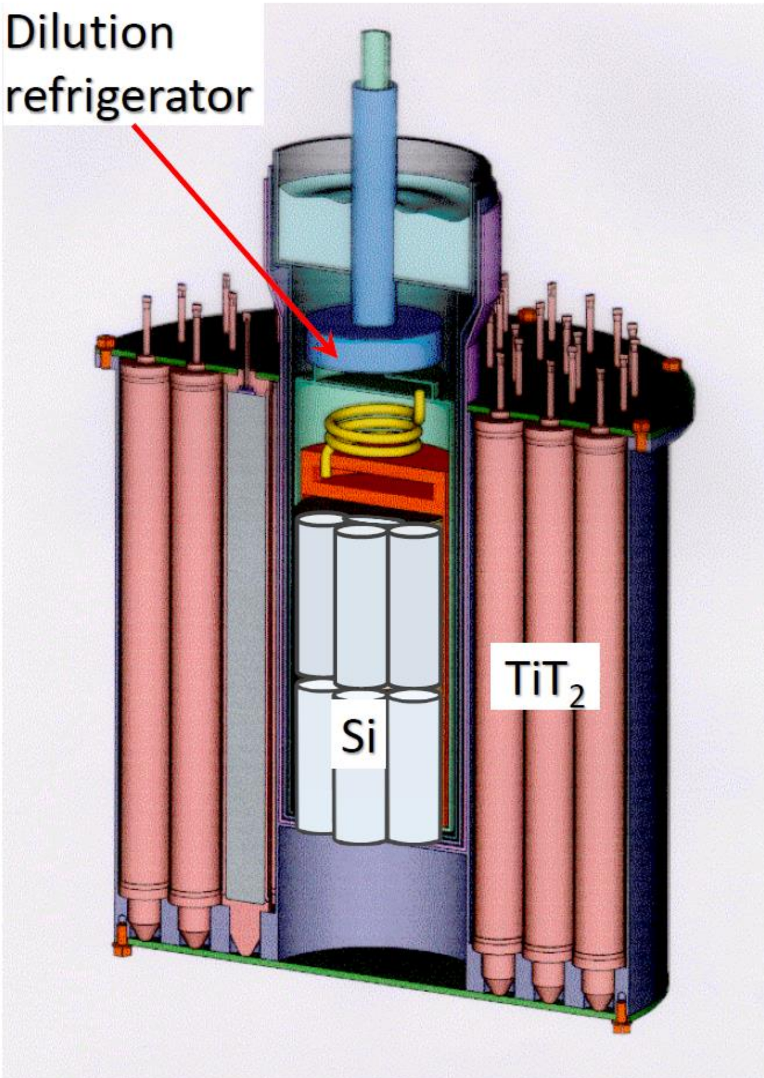
Differential cross sections for ionization of Si by tritium $\bar{\nu}_e$

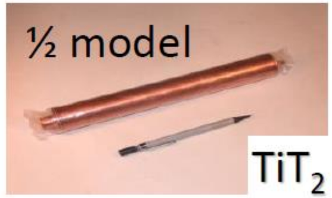
At small T values: $\frac{d\sigma_{SM}}{dT} \propto \mathrm{const}$, and $\frac{d\sigma_{(\mu)}}{dT} \propto \frac{\mu_\nu^2}{T}$



The detector's energy threshold is to be as low as possible

Si detector concept





Tritium neutrino source (1-4 kg)

- tubular elements with TiT_2

Silicon cryodetectors (T=10-50 mK)
(14-28)×125 cm³, M=2.9-5.7 kg

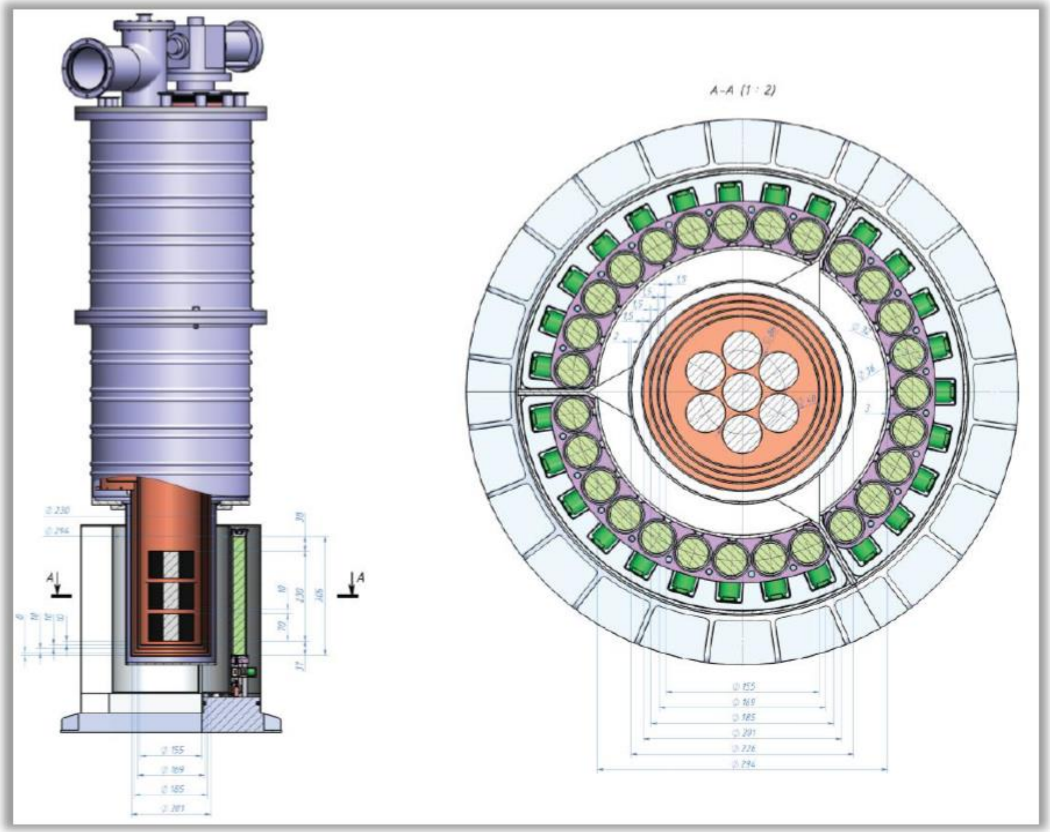
with TES or CEB mounted on each Si crystal

The Si detector with an ultra-low threshold E_{th}~10 eV or even E_{th}~1 eV owing to the Neganov-Trofimov-Luke effect *(heat amplification of ionization signal)*

B. Neganov and V. Trofimov, USSR patent no. 1037771, Otkrytia i Izobreteniya ***146*** *(1985) 215;*

P. N. Luke, J. Appl. Phys. ***64*** *(1988) 6858*.

Projected μ_ν-sensitivity of Si detector



Tritium mass is 1 kg (10 MCi)

$$\Delta\chi^2 = \chi^2 - \chi^2_{\min}$$

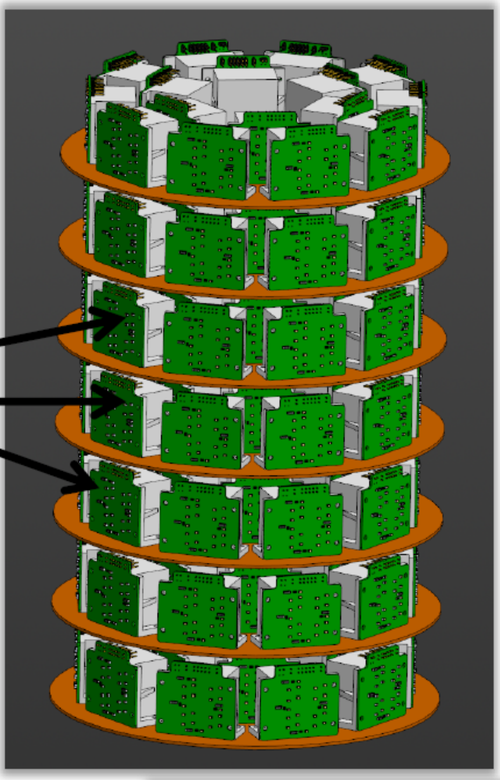
$$\chi^2 = \left(\frac{N_{SM} - N}{\sqrt{N_{SM}}}\right)^2$$

$$N = N_{SM} + N_{\mu_\nu}$$

1 year of taking data	**14 cylinders, 2.9 kg**		**21 cylinders, 4.3 kg**		**28 cylinders, 5.7 kg**	
	1 эВ	10 эВ	1 эВ	10 эВ	1 эВ	10 эВ
N_{SM}	7.96	7.94	11.52	11.49	14.61	14.57
μ_ν, $10^{-12}\mu_B$ 90% CL	1.76	2.03	1.61	1.85	1.51	1.74

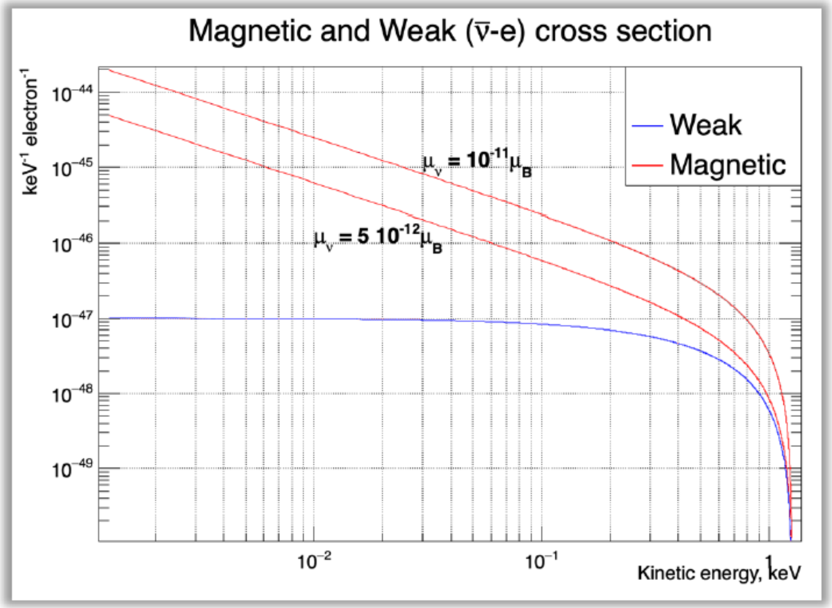
CsI(pure) detector concept

Detector assembly

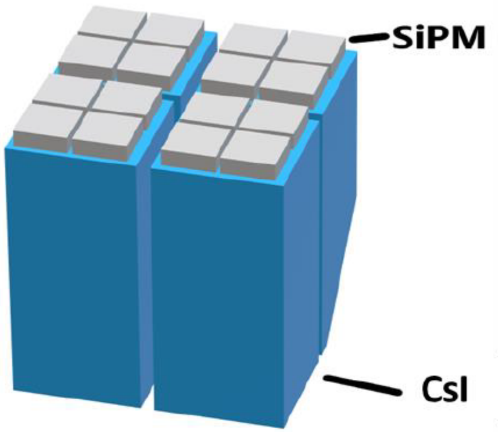


Layers of modules

*Abdurashitov, Vlasenko, Ivashkin, Silaeva, Sinev, Phys. Atom. Nuclei **85** (2022) 701*



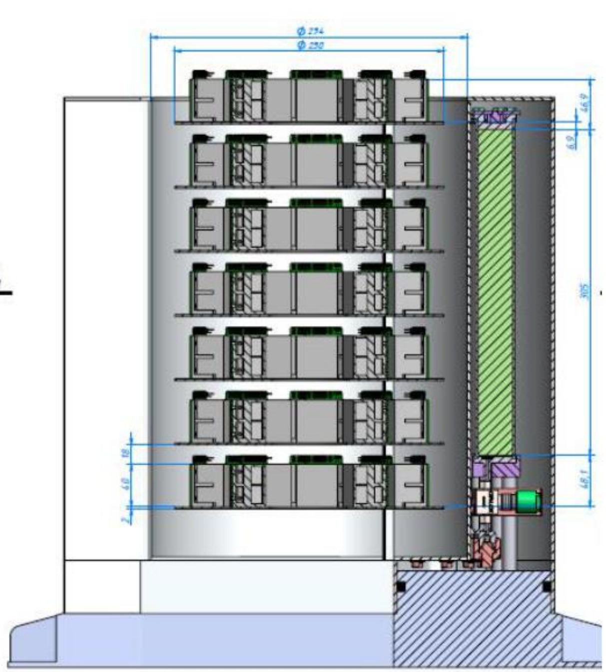
Detector module

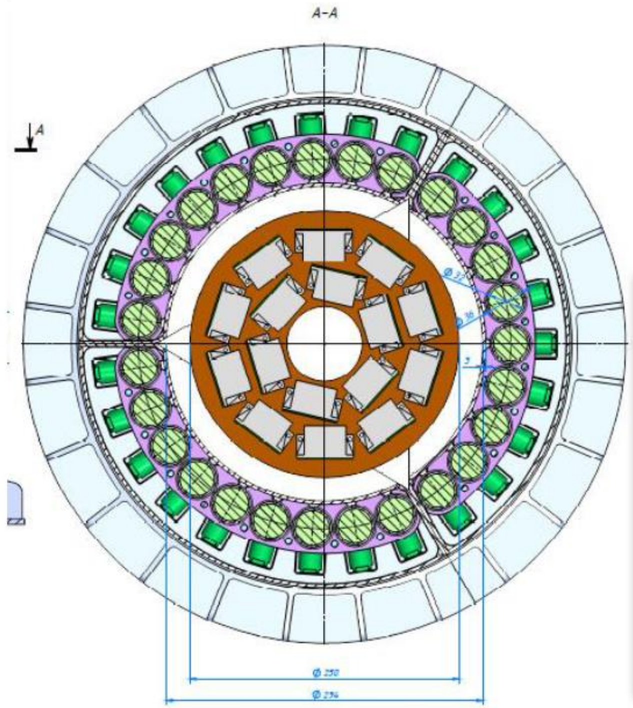


15x15x25 mm^3 CsI(pure) crystals at T=77 K, total mass is M=7.5-10.5 kg

- SiPM readout (two SiPMs per each crystal)
- Light collection at a level of ~30 photoelectrons/keV
- Energy threshold is E_{th}~100 eV

Projected μ_ν-sensitivity of CsI detector





Tritium mass is 1 kg (10 MCi)

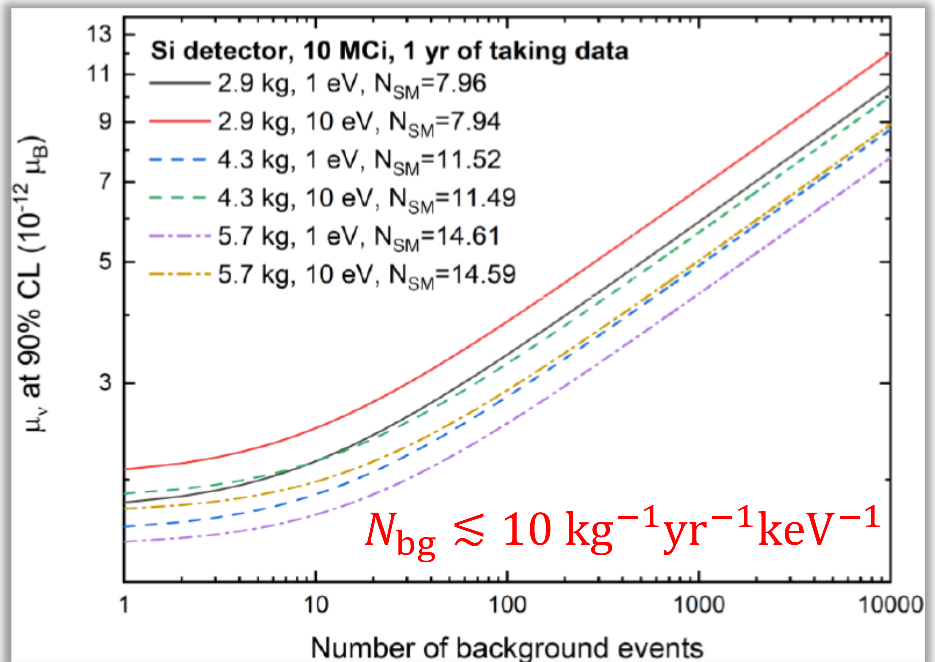
$$\Delta\chi^2 = \chi^2 - \chi^2_{\min}$$

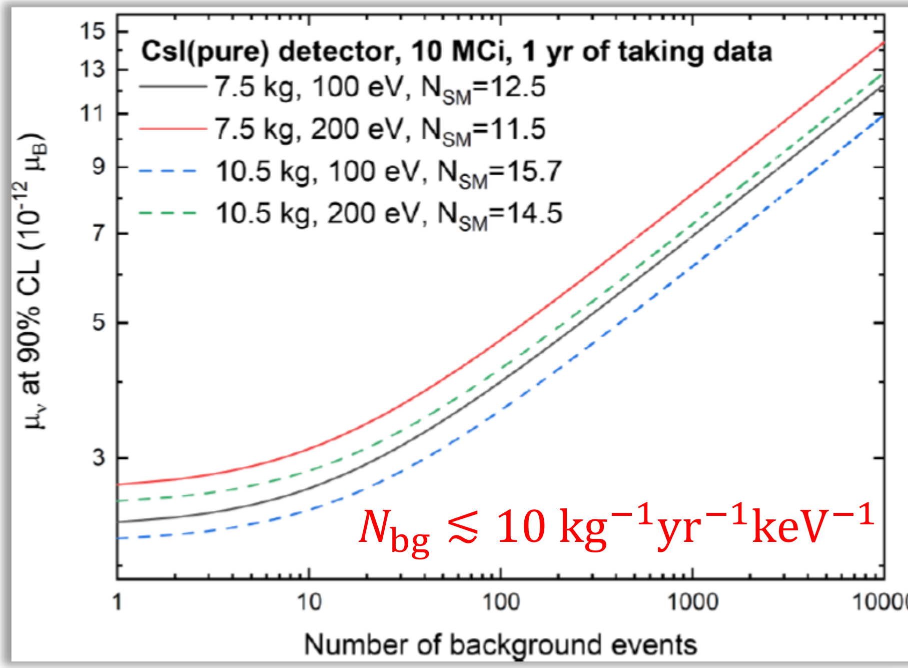
$$\chi^2 = \left(\frac{N_{SM} - N}{\sqrt{N_{SM}}}\right)^2$$

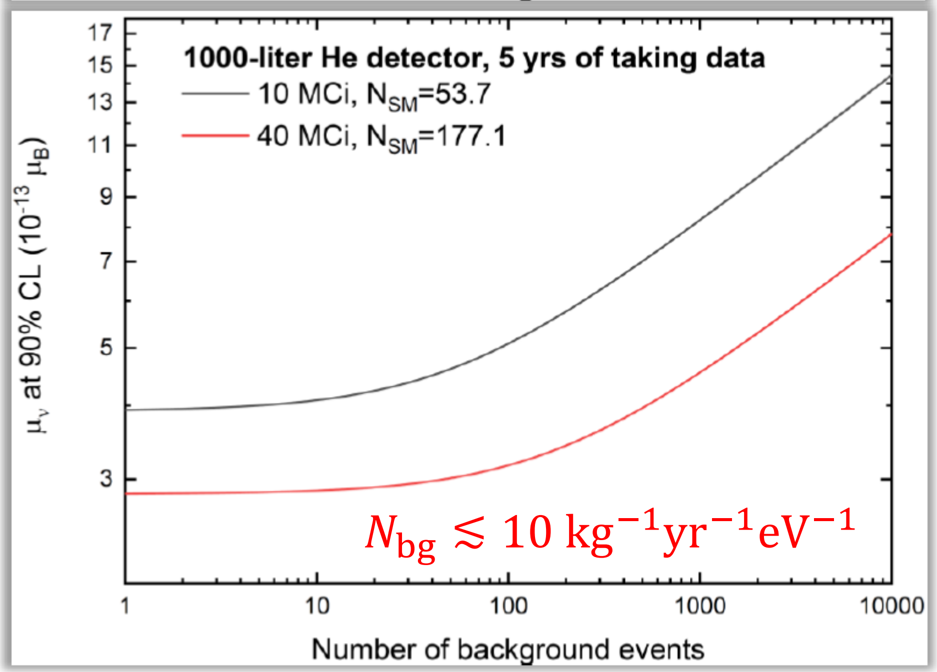
$$N = N_{SM} + N_{\mu_\nu}$$

1 year of taking data	**5 layers, 7.5 kg**				**7 layers, 10.5 kg**			
	100 эВ	200 эВ	300 эВ	400 эВ	100 эВ	200 эВ	300 эВ	400 эВ
N_{SM}	12.48	11.53	10.52	9.50	15.71	14.51	13.24	11.96
μ_ν, $10^{-12}\mu_B$ 90% CL	2.31	2.66	2.91	3.11	2.18	2.51	2.75	2.93

Required background conditions







To provide background conditions necessary to achieve the world-leading sensitivities to the neutrino magnetic moment, some measurements may be performed in the Baksan Neutrino Observatory

Summary and outlook for SATURNE

Sarov tritium neutrino experiment (SATURNE) aims at

(i) first ever observation of **CEνAS** to test SM neutrino interactions at unprecedentedly low energies

(ii) search for **neutrino magnetic moment**

High-intensity tritium neutrino source is being prepared

- at least **1 kg, 10 MCi** (possibly up to **4 kg, 40 MCi**)

He II detector is being developed

- observation of **CEνAS at 5σ (2030-2035)**
- sensitivity to $\mu_\nu \sim$**(3-4)x10^{-13} μ_B at 90% CL**

Si detector is being developed

- sensitivity to $\mu_\nu \sim$**(1.5-2.0)x10^{-12} μ_B at 90% CL (2029-2030)**

CsI(pure) detector is being developed

- sensitivity to $\mu_\nu \sim$**(2-3)x10^{-12} μ_B at 90% CL (2028-2029)**

НЦФМ
НАЦИОНАЛЬНЫЙ ЦЕНТР
ФИЗИКИ И МАТЕМАТИКИ

The SATURNE Collaboration

RUSSIAN FEDERAL NUCLEAR CENTER
ALL-RUSSIAN RESEARCH INSTITUTE OF EXPERIMENTAL PHYSICS

PRODUCTION
ASSOCIATION
MAYAK

NIZHNY NOVGOROD STATE
TECHNICAL UNIVERSITY
N.A. R.E. ALEKSEEV

LOMONOSOV MOSCOW
STATE UNIVERSITY

LOBACHEVSKY
UNIVERSITY

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Physical-
Technical
Institute

Future Neutrino Experiments

Neutrino Unbound

https://www.nu.to.infn.it

- **ANNIE** Neutrino Interactions (Home, INSPIRE, Wikipedia) References
- **DUNE** Accelerator LBL Oscillations, Atmospheric and Supernova Neutrinos, Proton Decay (Home, INSPIRE, WikipediA) References
- **ECHO** Electron Neutrino Mass (Home, INSPIRE) References
- **ESSnuSB** Accelerator LBL Oscillations (Home, INSPIRE) References
- **GRAND** High-Energy Astrophysical Neutrinos (Home, INSPIRE) References
- **HOLMES** Electron Neutrino Mass (Home, INSPIRE) References
- **HUNT** High-Energy Astrophysical Neutrinos (Home, INSPIRE) References
- **Hyper-Kamiokande** Accelerator LBL Oscillations, Atmospheric and Supernova Neutrinos, Proton Decay (Home, INSPIRE, WikipediA) References
- **JSNS^2** Accelerator SBL Oscillations, Experiment (Home, INSPIRE) References
- **JNE** Solar, Geo and Supernova Neutrinos (Home, INSPIRE) References
- **JUNO** Reactor LBL Oscillations, Atmospheric, Solar, Geo Neutrinos (Home, INSPIRE, WikipediA) References
- **LEGEND** Neutrinoless Double Beta Decay (^{76}Ge) (Home) References
- **P-ONE** High-Energy Astrophysical Neutrinos (Home, INSPIRE) References
- **SATURNE** Coherent Elastic Neutrino-Atom Scattering (INSPIRE) References
- **SBN** Accelerator SBL Oscillations, and Experiment (Home, INSPIRE) References
- **TRIDENT** High-Energy Astrophysical Neutrinos (Home, INSPIRE) References
- **WATCHMAN** Reactor Anti-Neutrino Monitor (Home, INSPIRE) References

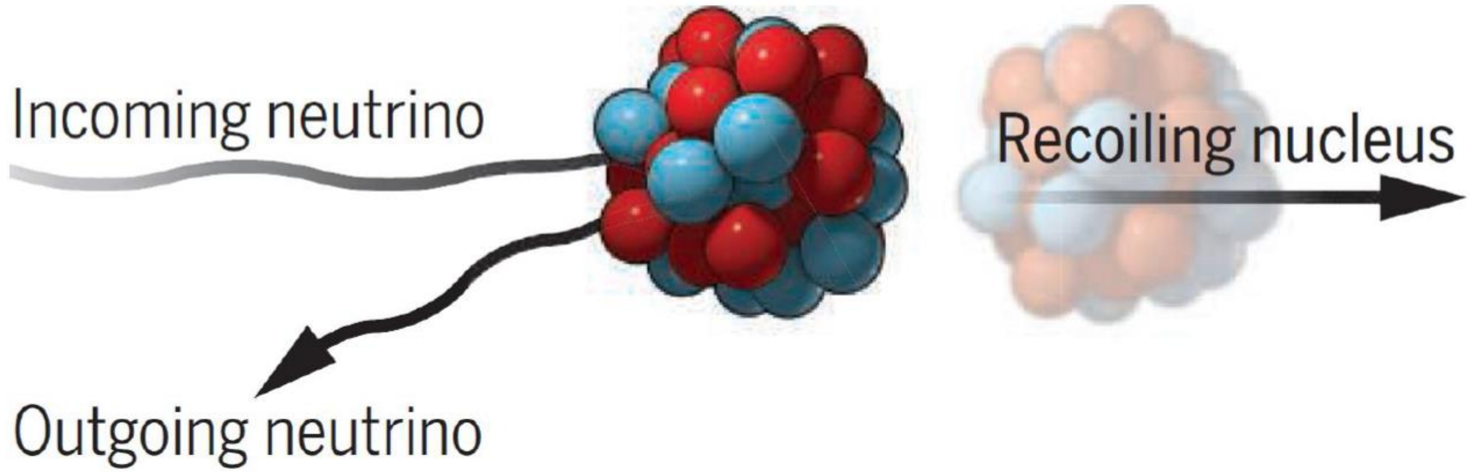
Единственный CE*ν*AS эксперимент
+
уникальный потенциал для поиска ЭМ свойств нейтрино

Спасибо за внимание!

BACKUP

CEνNS: Coherent Elastic ν-Nucleus Scattering

Predicted by *D. Z. Freedman, PRD **9** (1974) 1389;*
*V. B. Kopeliovich & L. L. Frankfurt, ZhETF Pis. Red. **19**, No. 4 (1974) 236*



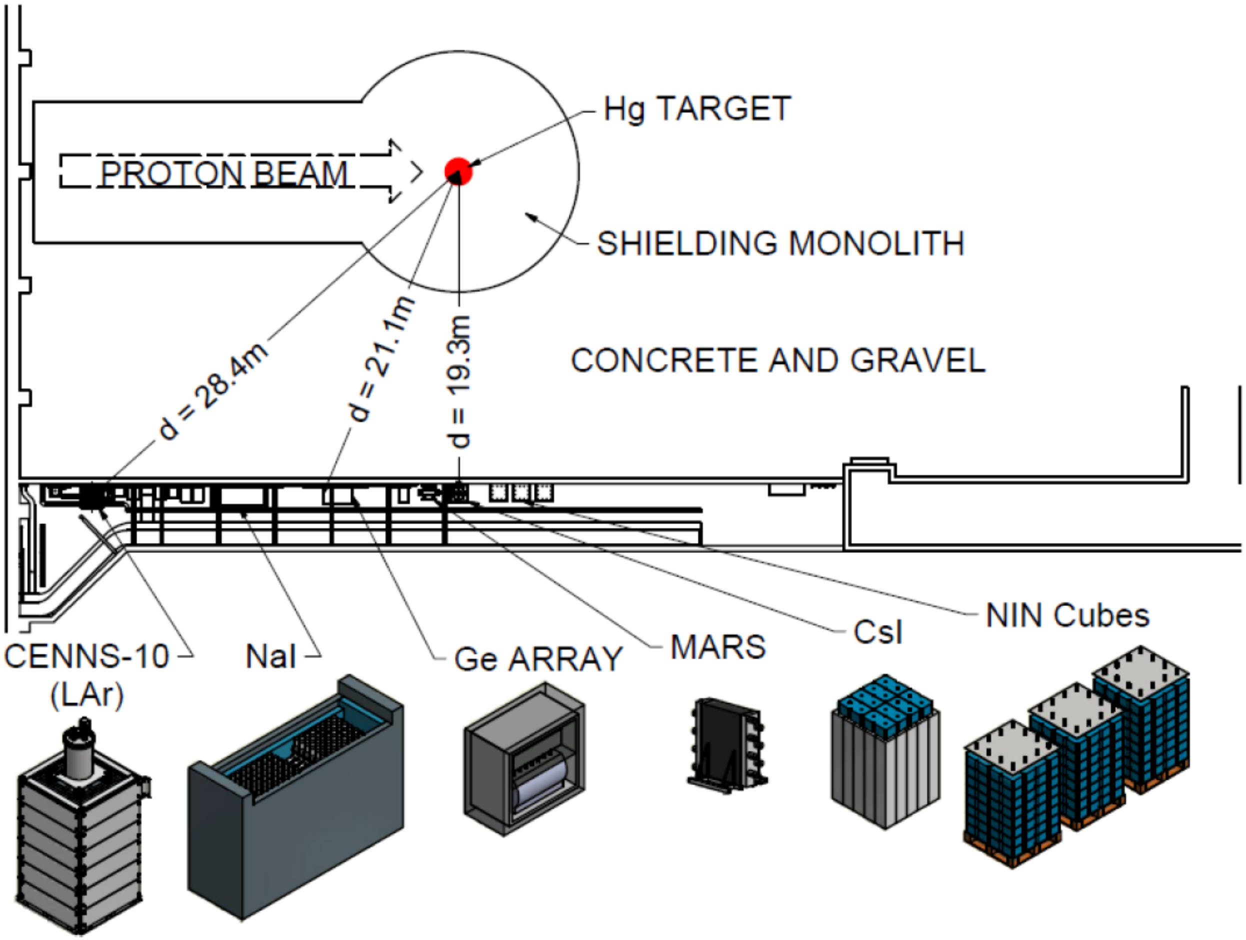
Нейтрино рассеивается на ядре атома. При этом внутренняя энергия ядра не меняется, а амплитуды рассеяния на отдельных нуклонах ядра складываются когерентным образом. В Стандартной модели:

$$\frac{d\sigma_{\mathrm{CE}\nu NS}}{dT} \underset{\sim}{\propto} N^2$$

The first ever observation of CEνNS

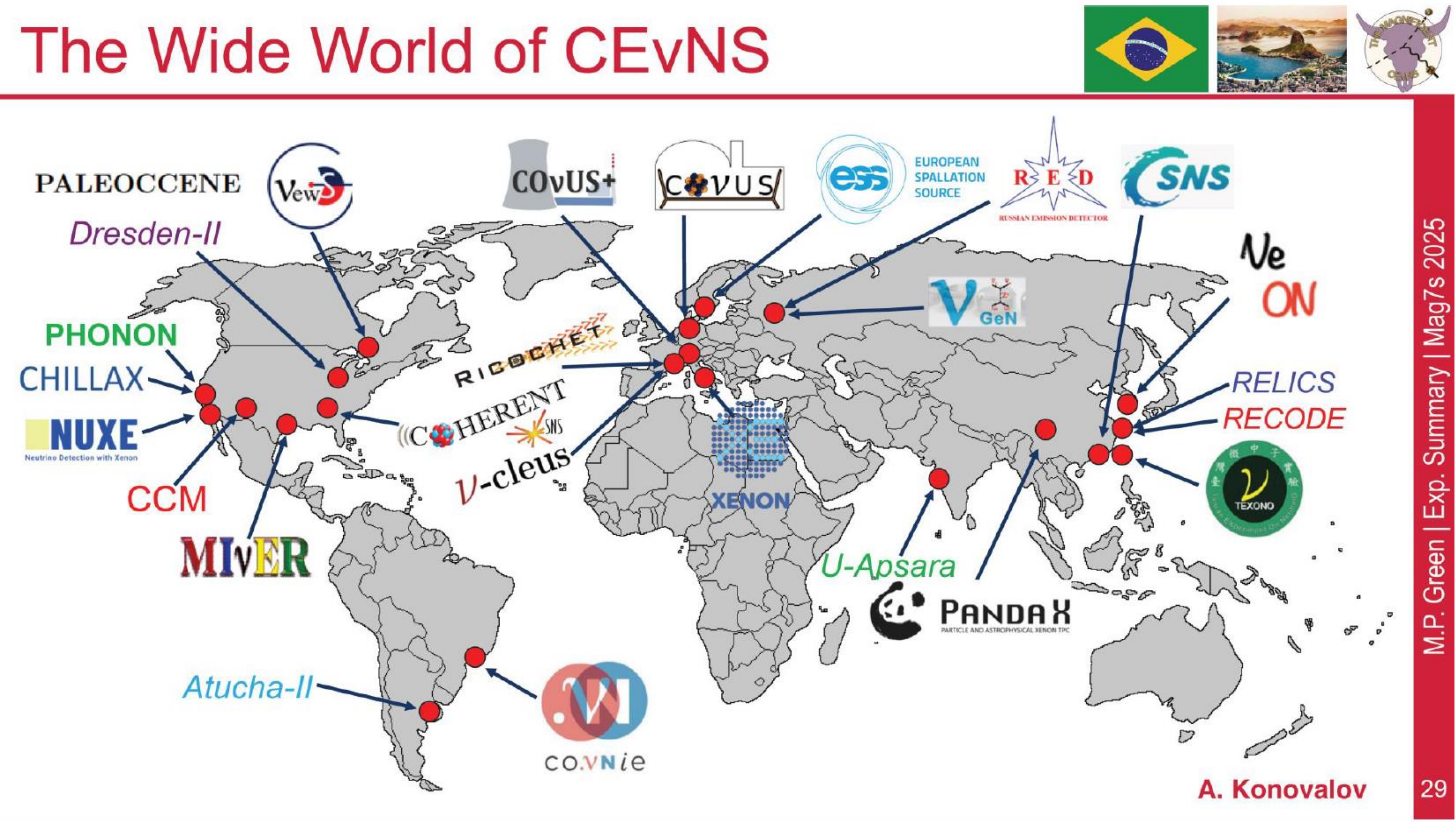
*D. Akimov et al. (COHERENT Collab.), Science **357** (2017) 1123*

The COHERENT experiment (SNS, Oak Ridge, Tennessee)



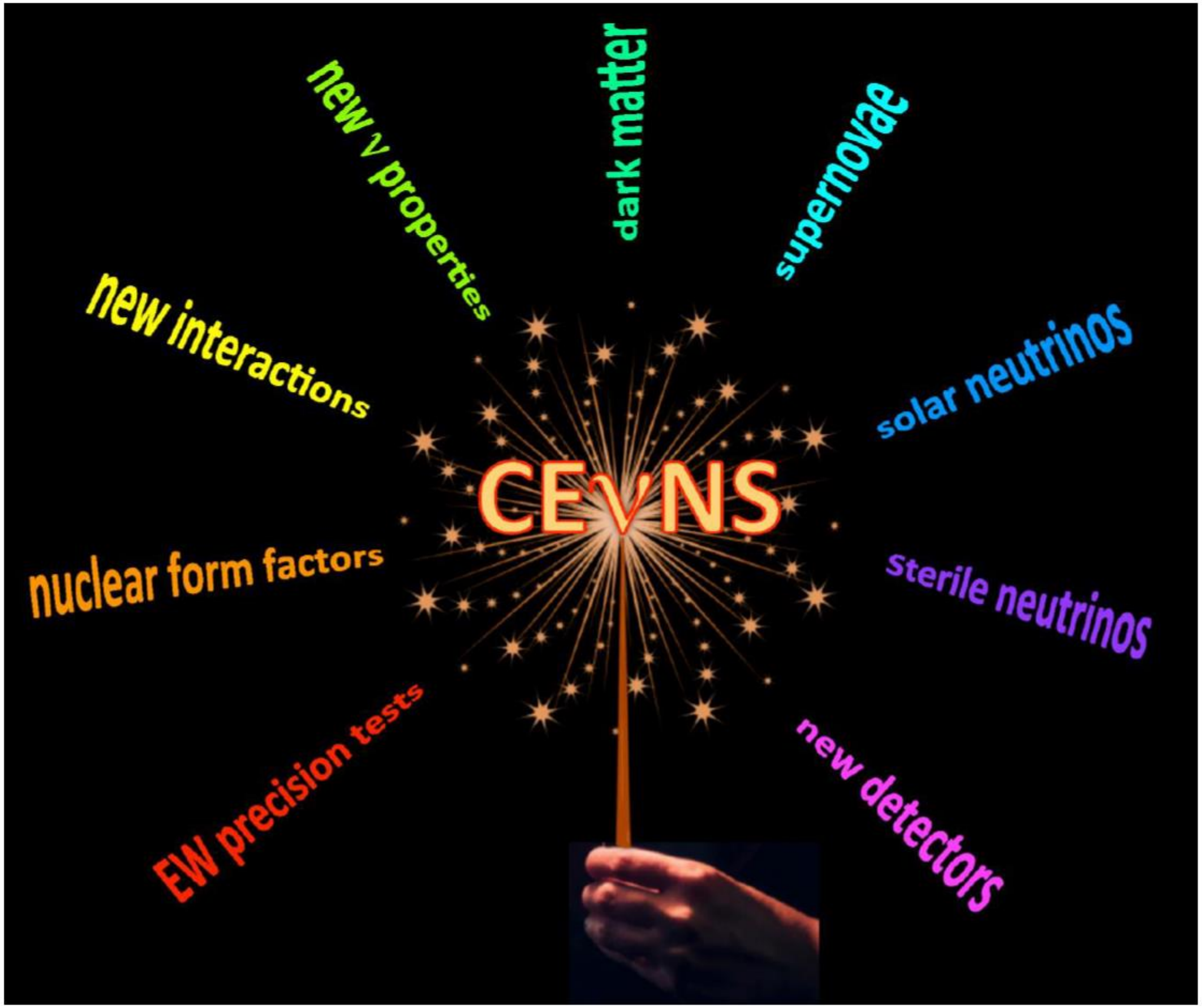
14.6 kg CsI
scintillating crystal

Global map of current CEνNS experiments



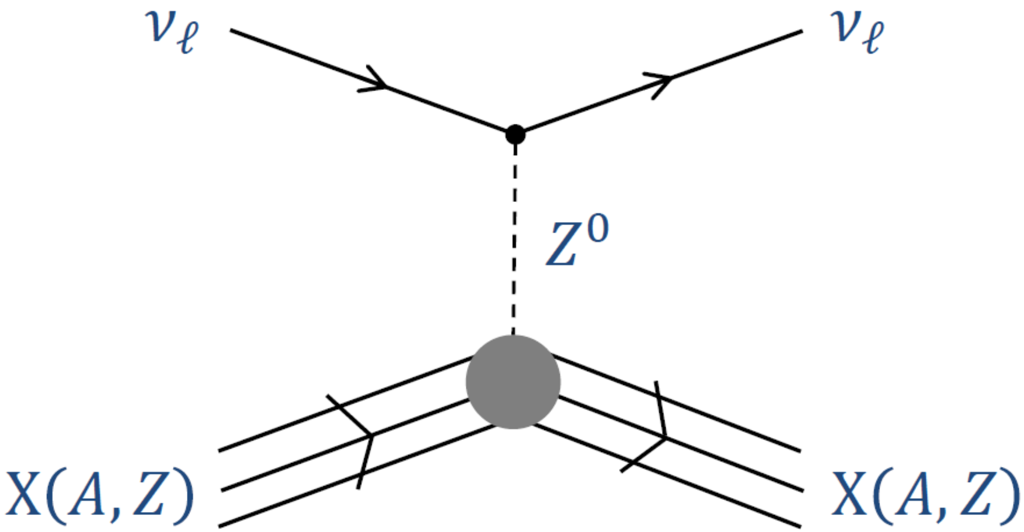
[M.P. Green, Magnificent CEνNS 2025]

CEνNS potential



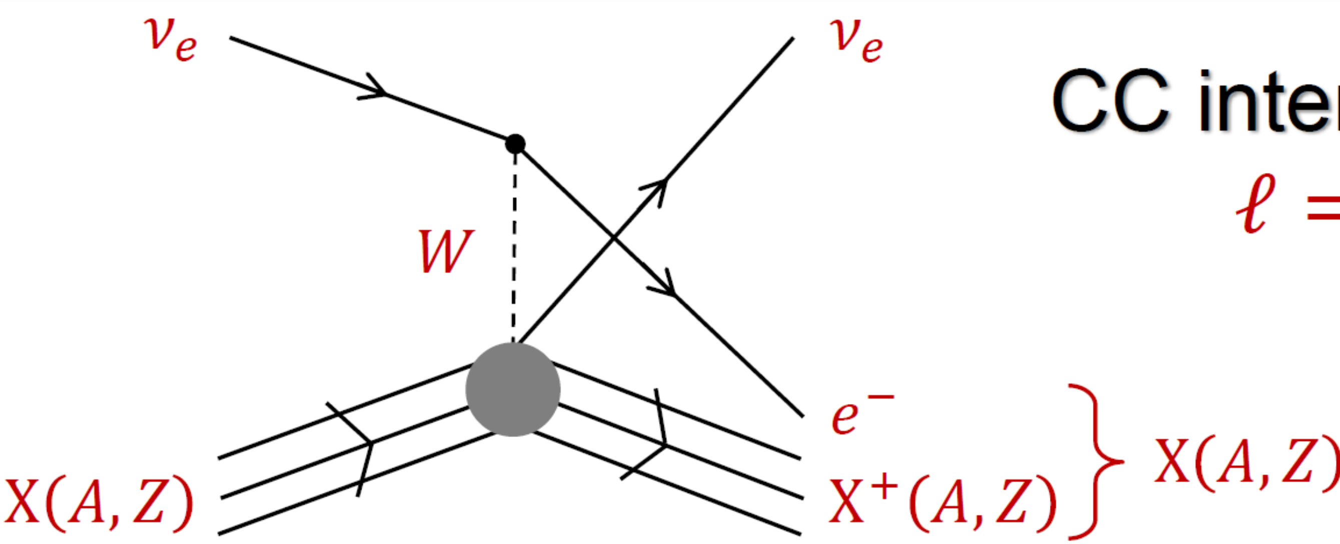
[E. Lisi, Neutrino 2018]

Elastic neutrino-atom scattering in SM



NC interaction

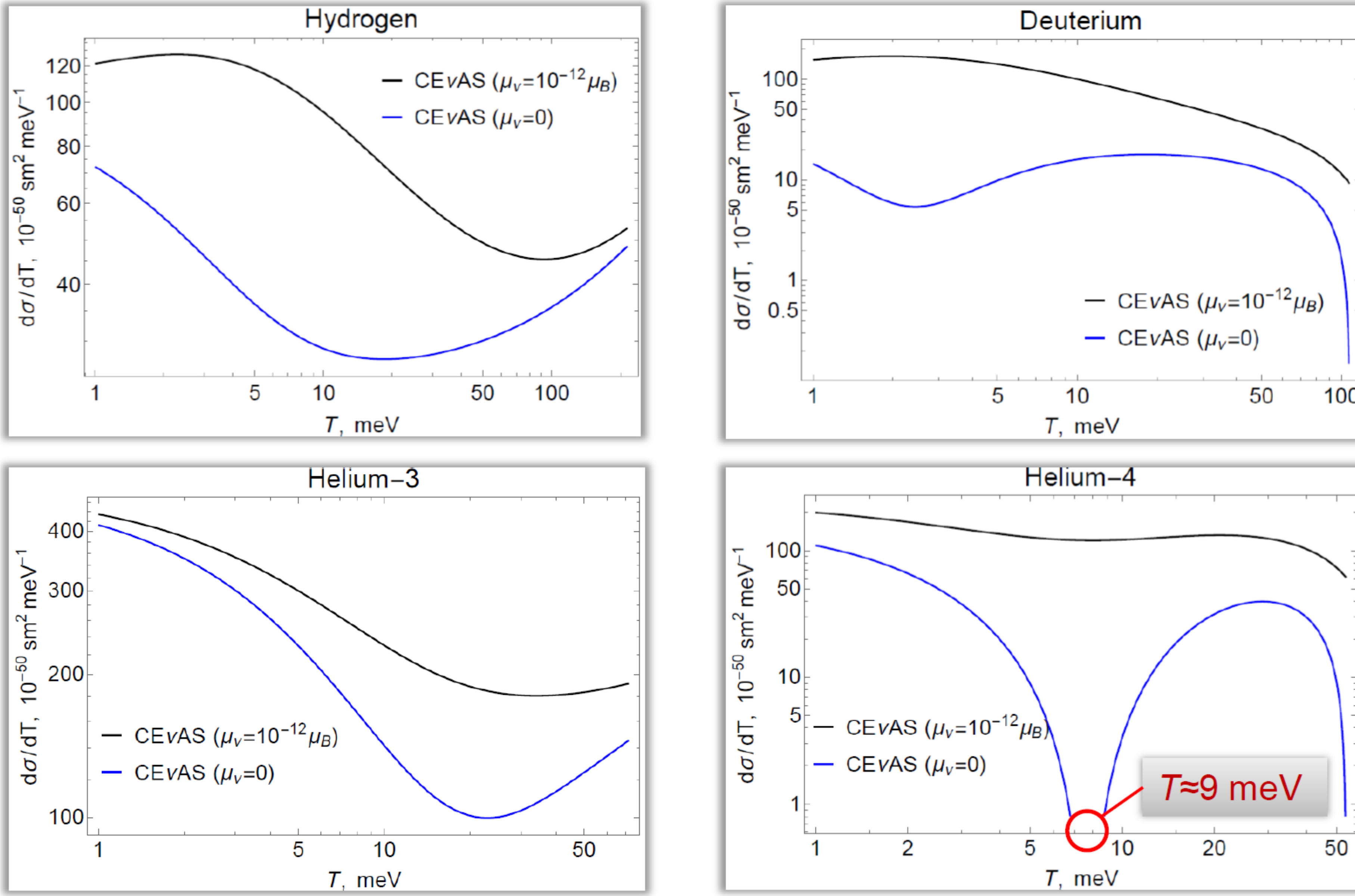
$\ell = e, \mu, \tau$



CC interaction

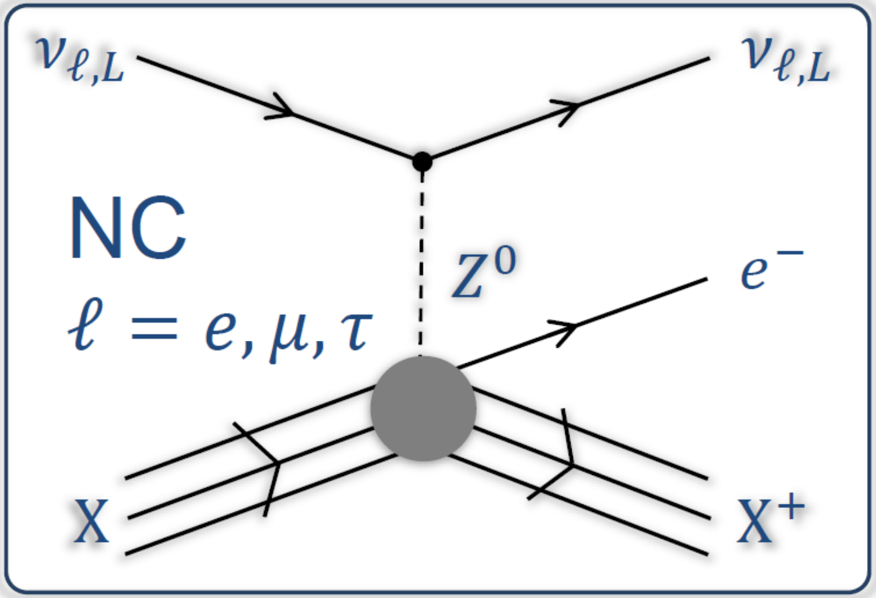
$\ell = e$

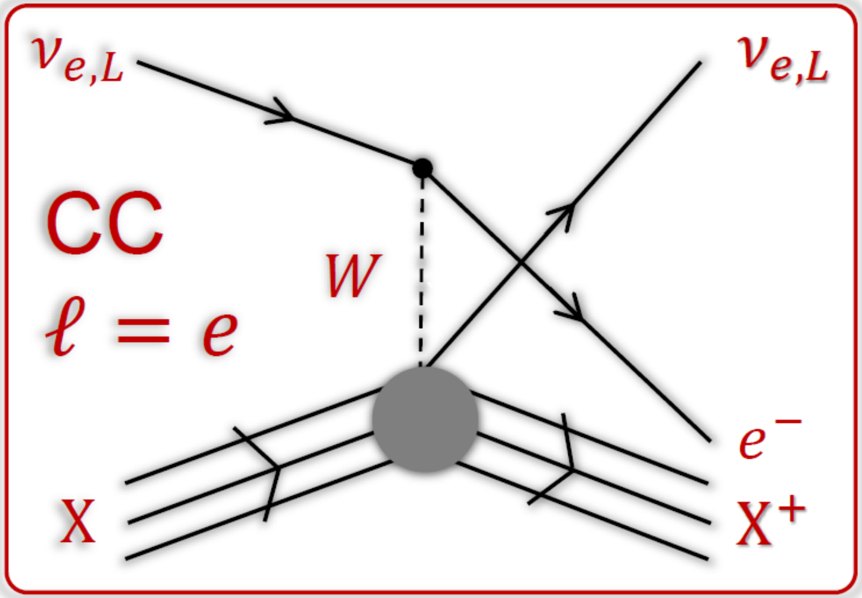
Effect of μ_ν on CEνAS cross sections for $\bar{\nu}_e$ with E_ν=10 keV

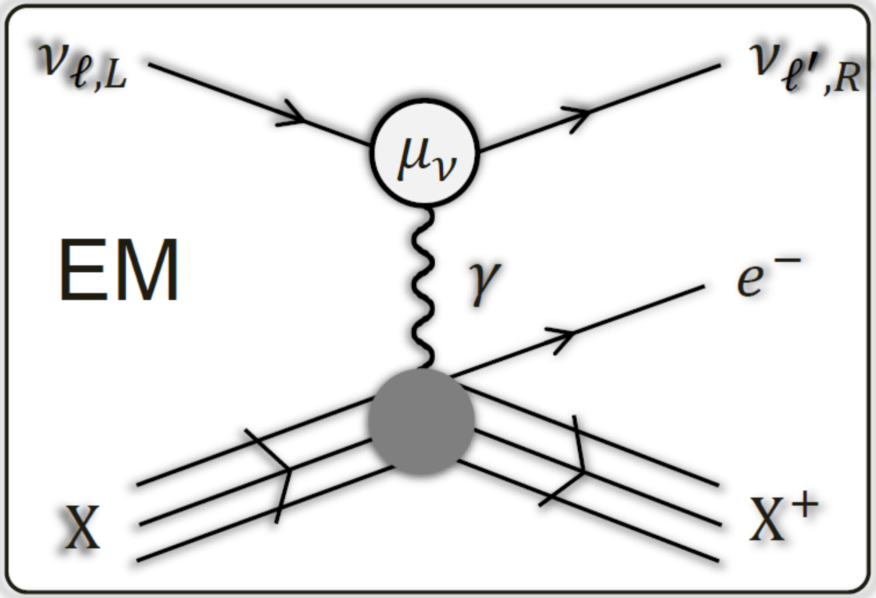


G. Donchenko, K. Kouzakov, A. Studenikin, JETP Lett. **117**, 879 (2023)

Ionizing neutrino-atom interactions

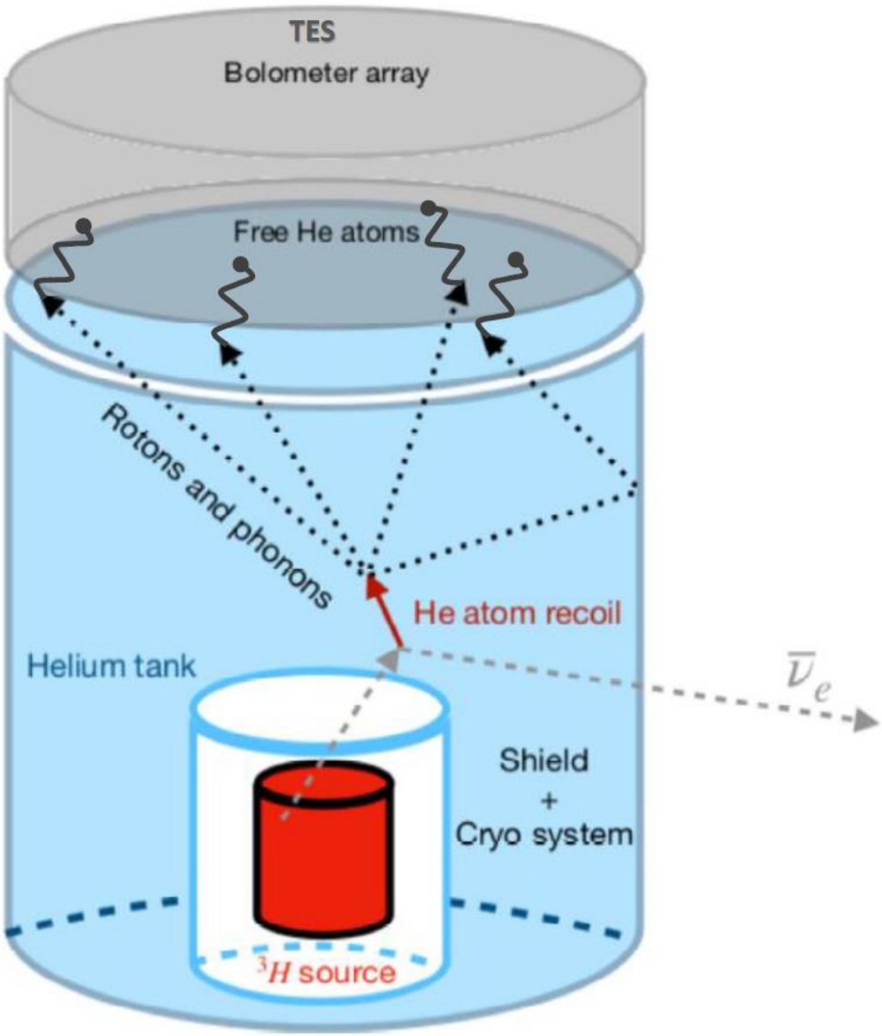






Unlike the SM weak NC and CC interactions, the μ_ν interaction flips the neutrino helicity ($L \to R$) and can change the neutrino flavor if the transition magnetic moment ($\ell \neq \ell'$) is nonzero.

Detection method to study CEνAS



Tritium neutrino

↓

Elastic neutrino-atom scattering

↓

He atom recoil (1-185 meV)

↓

Phonons and rotons (0.7-1.2 meV per q. p.)

↓

Quantum evaporation (0.6 meV per atom)

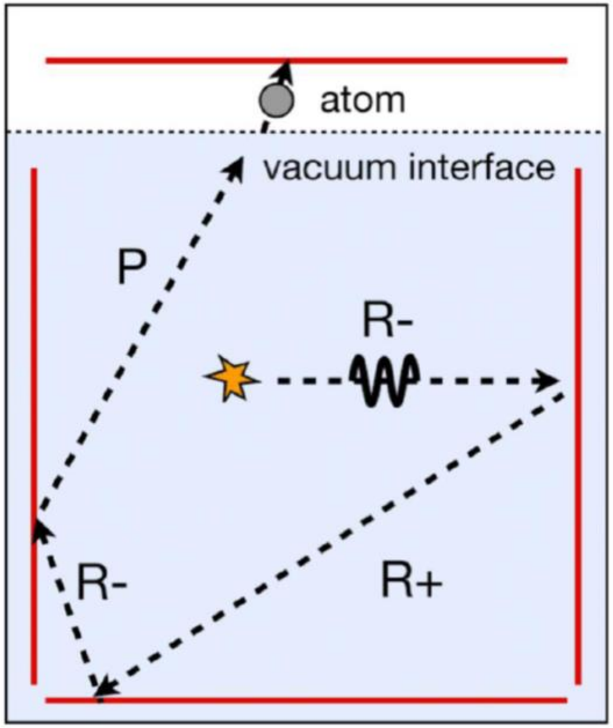
↓

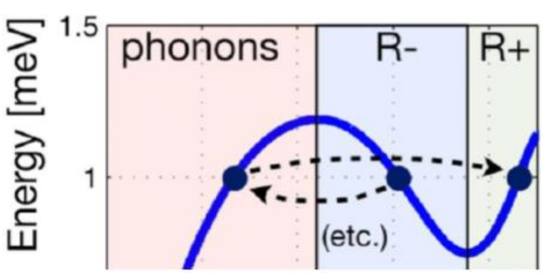
Adsorption onto TES (6-50 meV per adatom)

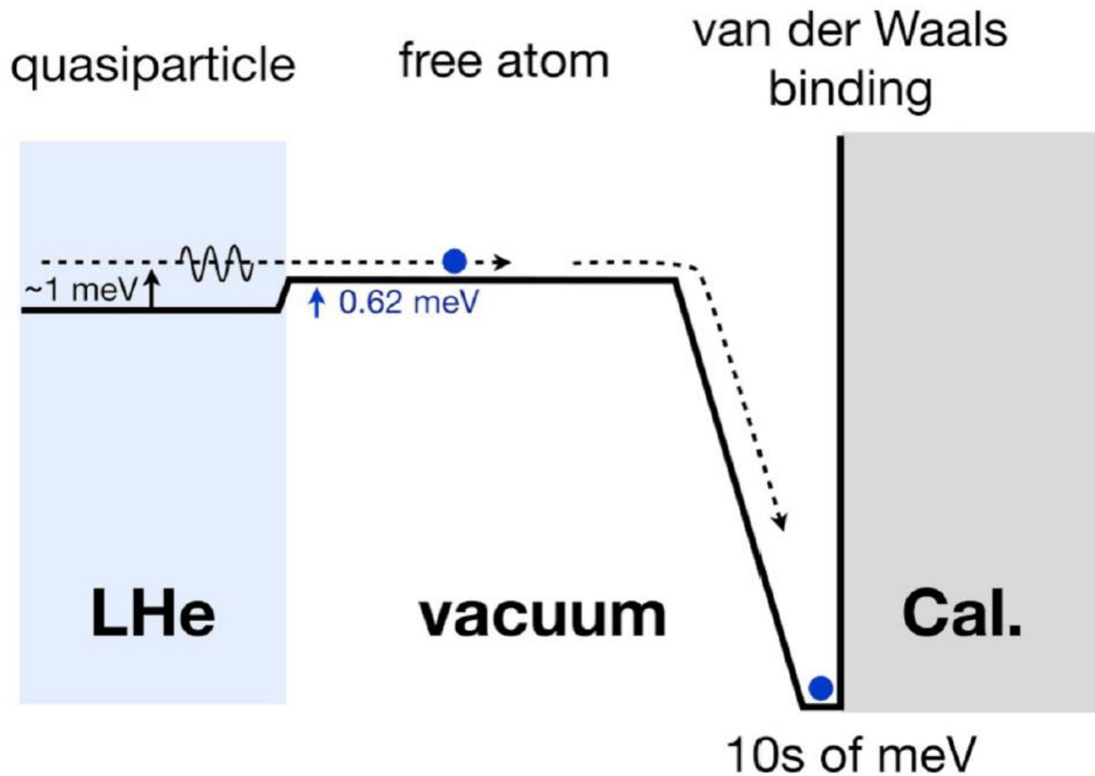
↓

Signal

Quasiparticle readout: Quantum evaporation of He atom







1 meV roton energy becomes up to 40 meV observable

- **x** 40 amplification
- Graphene-fluorine surface

[D. McKinsey, SNOLAB Workshop 2021]

The test He II cell for TRITON 200

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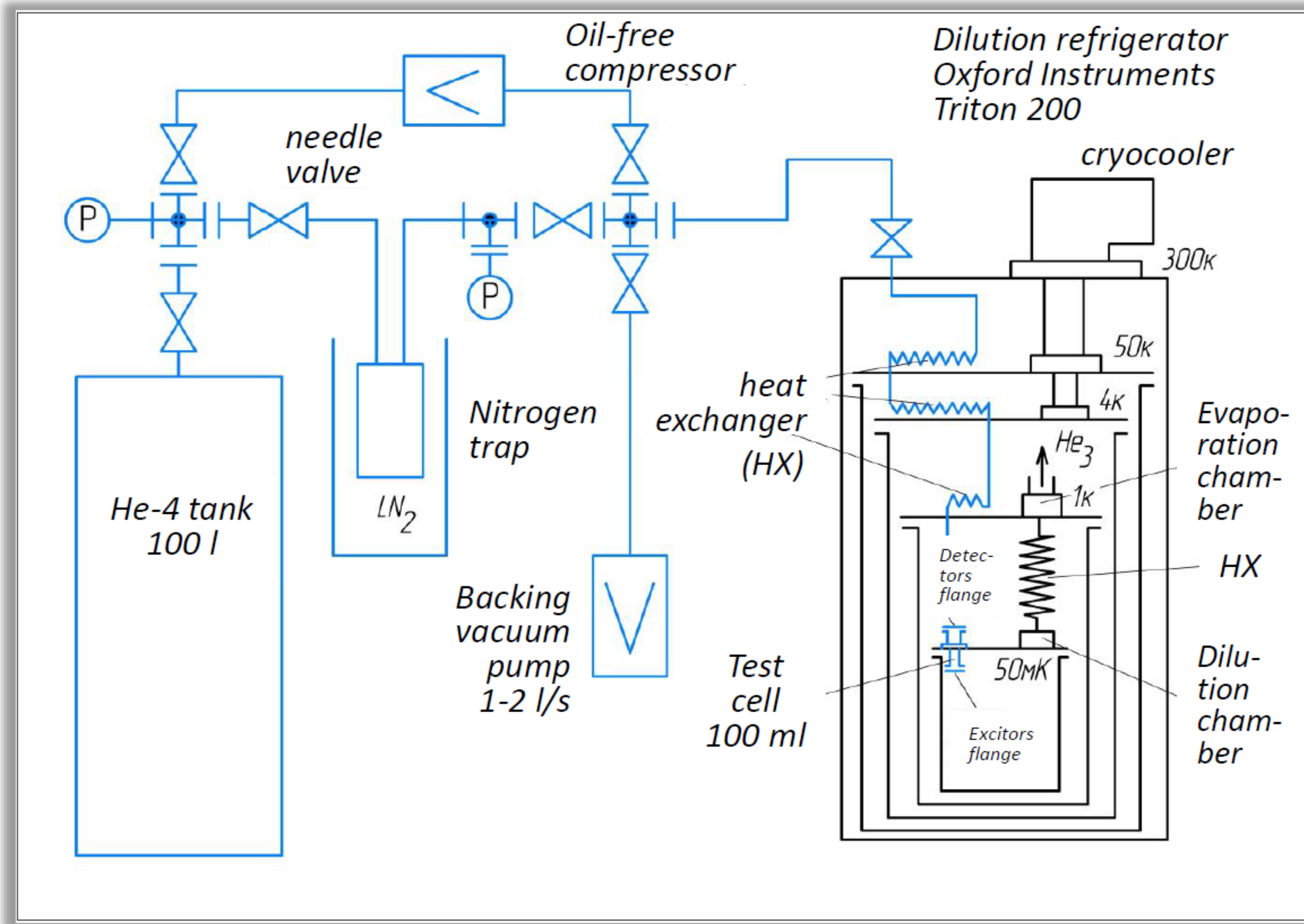
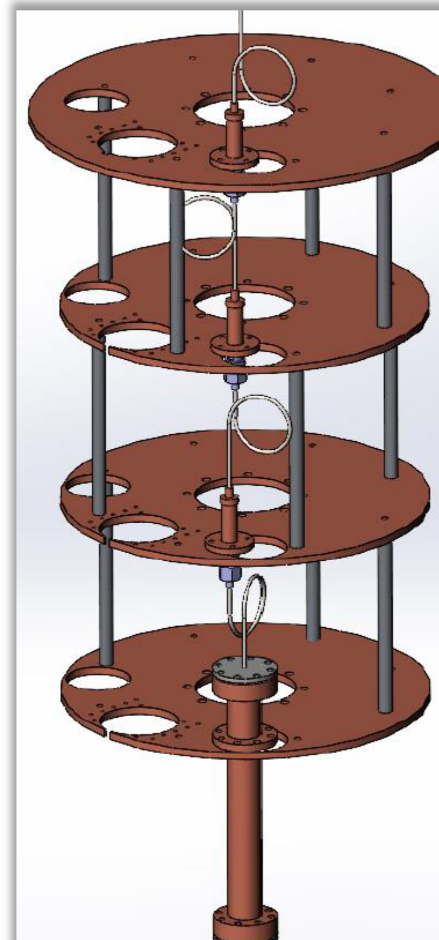
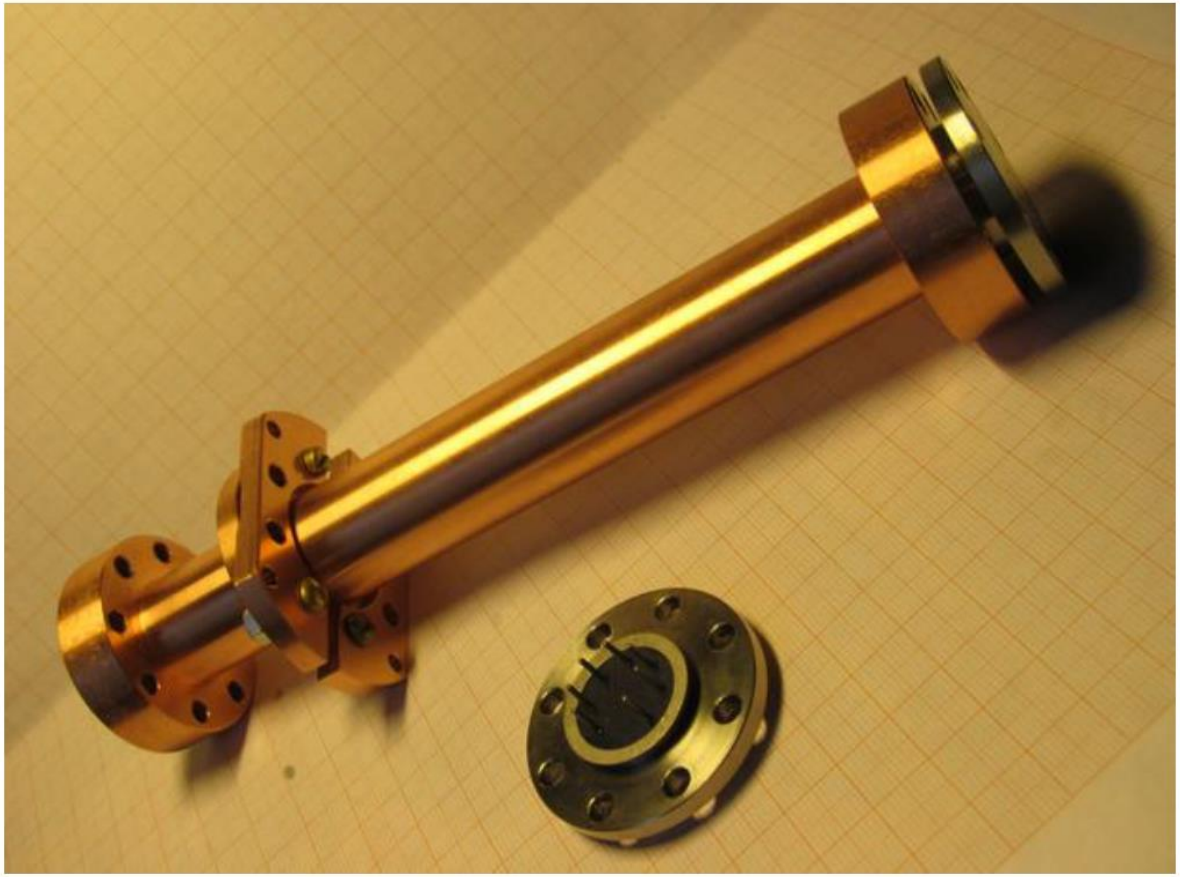




Purpose: To test the possibility of (i) generation of various excitations in helium (phonons, rotons, scintillations) by various controlled methods (thermal, mechanical, irradiation with various particles) and (ii) registration of these excitations by microcalorimetric detectors of various types

Тестовая ампула объемом 50 мл из бескислородной меди Моб

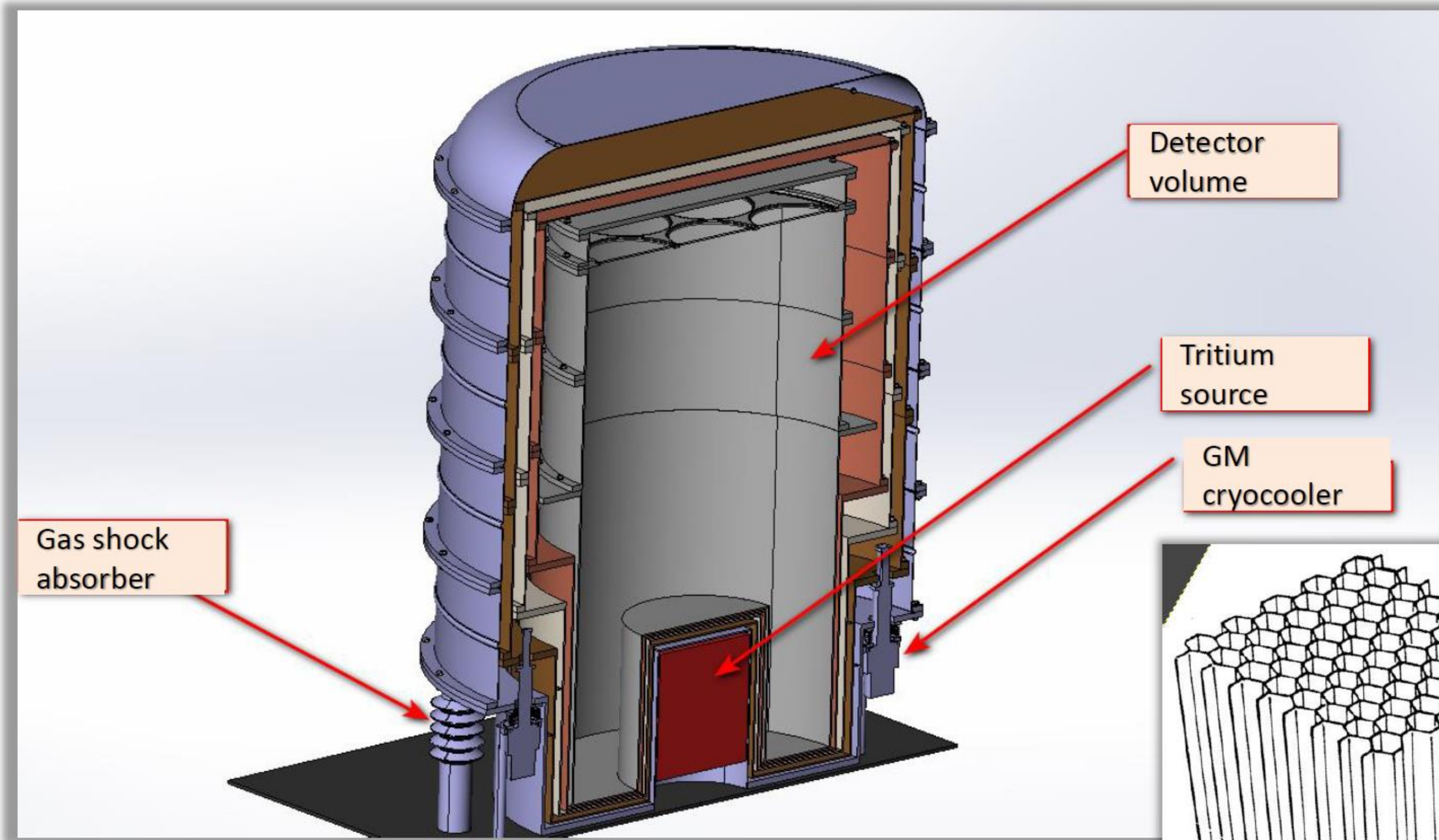


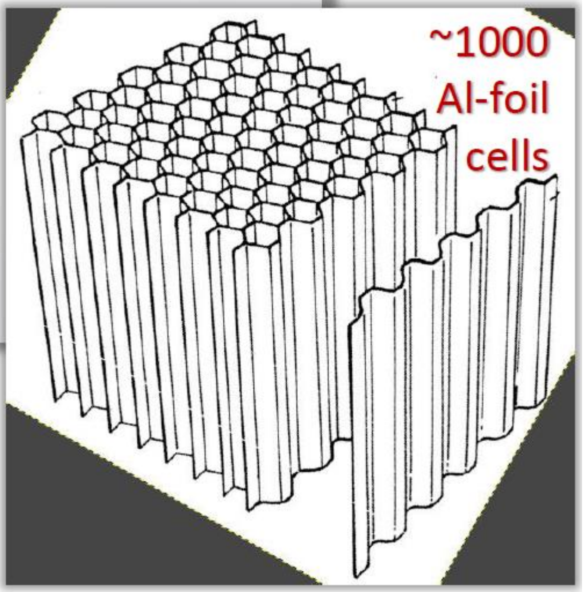
Ячейка имеет цилиндрическую геометрию и размещается вертикально, а верхний и нижний торцы закрываются фланцами из нержавеющей стали с индиевым уплотнением, через которые осуществляются электровводы генераторов и детекторов.

Внутренняя поверхность ампулы отполирована для увеличения коэффициента отражения ротонов от стенки.

Discrimination of background events

General view of the cryostat

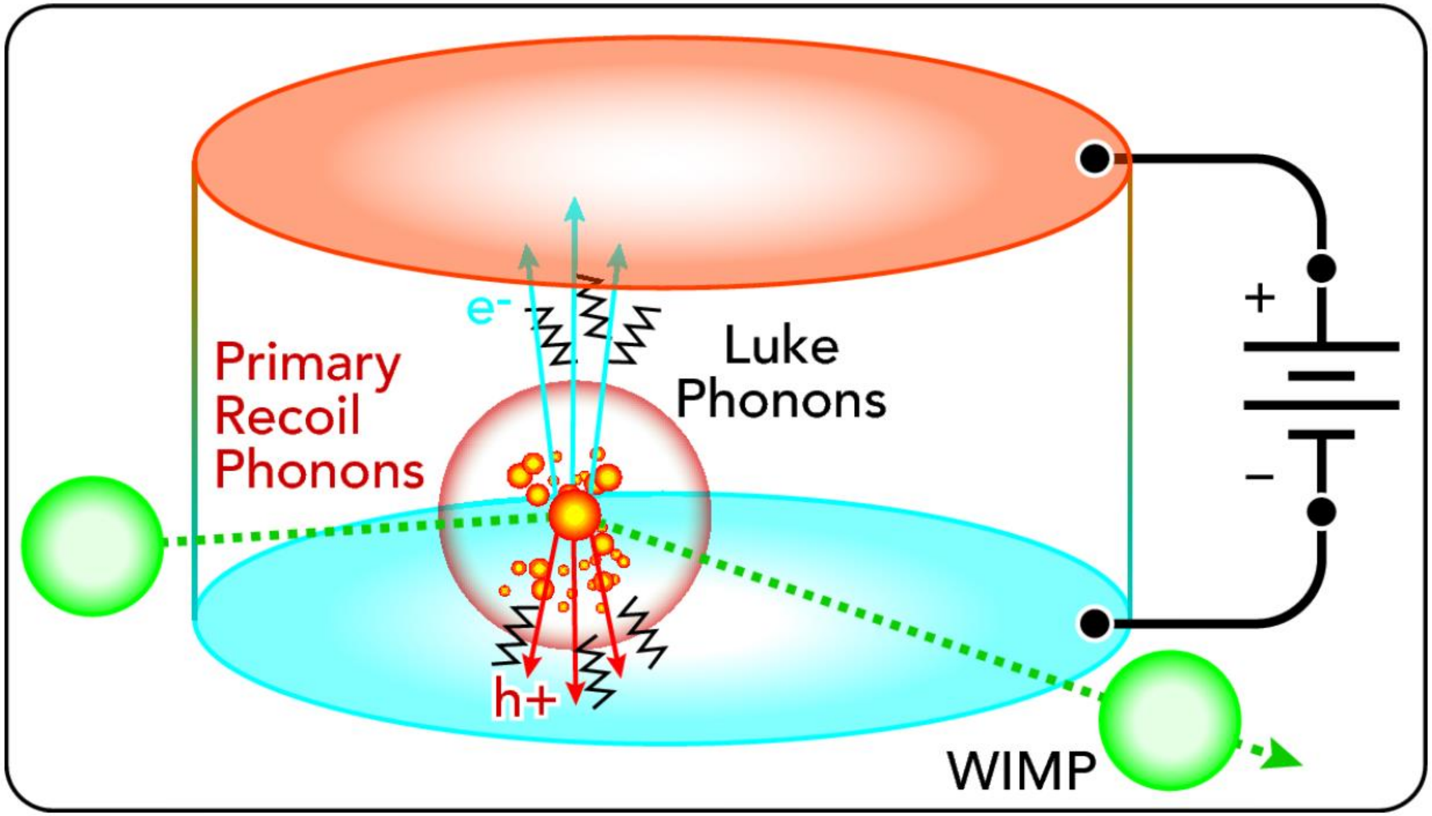




Segmentation of the He detector is the key to background discrimination

Neganov-Trofimov-Luke effect

Phonon amplification of ionization signal



Observed Phonon Energy = $E_{Recoil} + E_{NTL}$

[B. von Krosigk (on behalf of the SuperCDMS Collaboration), IDM2018]